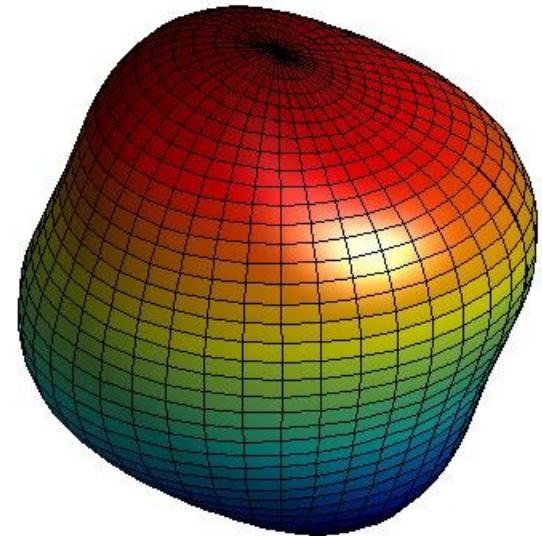


Stellar pulsations: chaotic or quasiperiodic

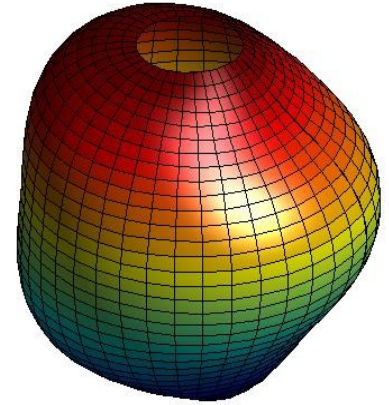
Viktor Votruba

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² Přírodovědecká fakulta MU Brno, ÚTFA



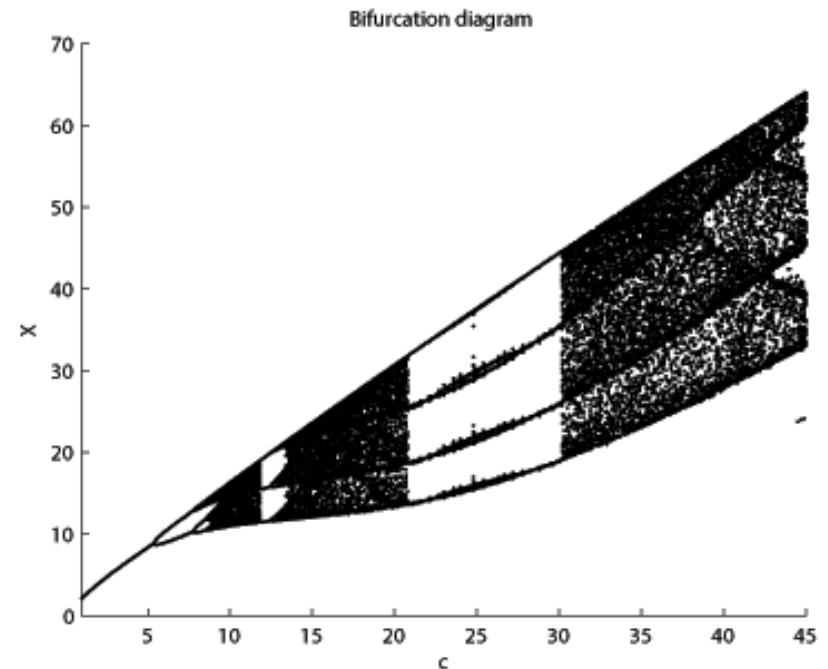
Deterministic chaos and pulsation



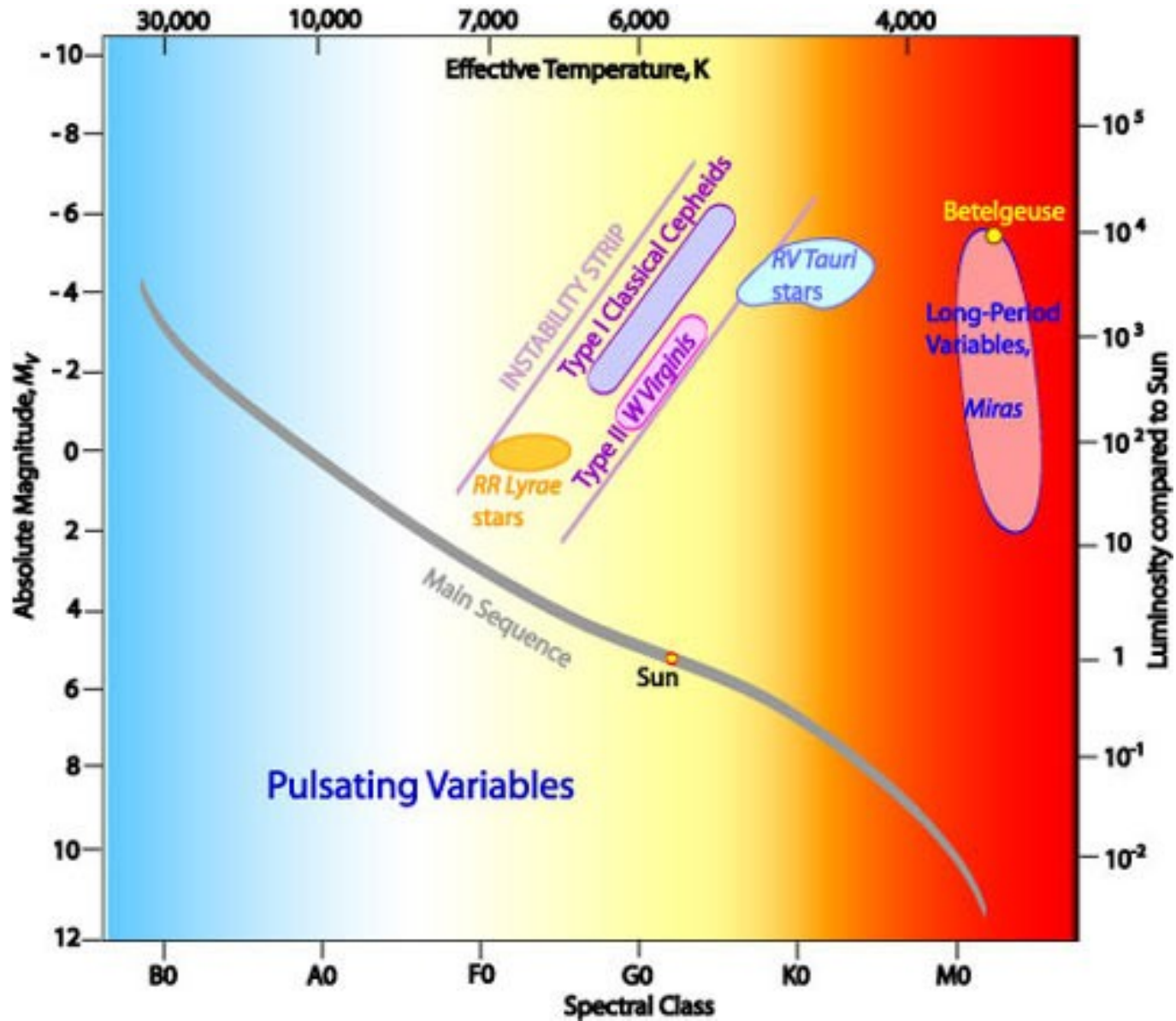
- Pulsating stars changing luminosity due to the radial and nonradial pulsation
- Basic mechanism of pulsation – κ mechanism
- Detection of regular or **irregular** photometric and spectral variability
- Recorded time - series are nonequidistant and very often with poor sampling and strong noise.
- Can be pulsations chaotic and can we find them ?

Low-dimensional deterministic chaos

- Set of coupled nonlinear ODE
- Strong sensitivity on initial conditions
- Phase portrait have fractal structure, strange attractor
- Change of driving parameters can lead to different regime



Pulsating stars



1D model of pulsations

Tanaka (1988)

Linear

$$\begin{aligned}\frac{dx}{dt} &= y \\ \frac{dy}{dt} &= \alpha x + z \\ \frac{dz}{dt} &= -\beta \frac{dx}{dt}\end{aligned}$$

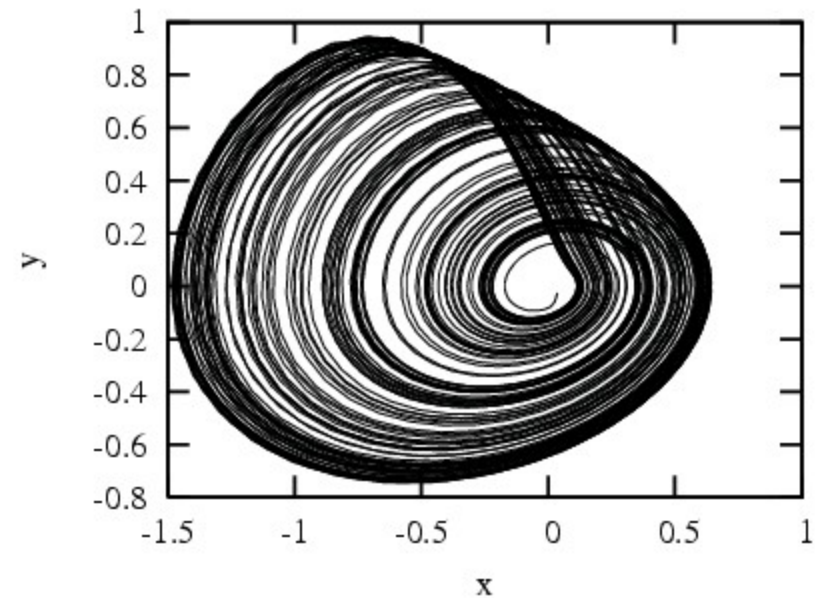
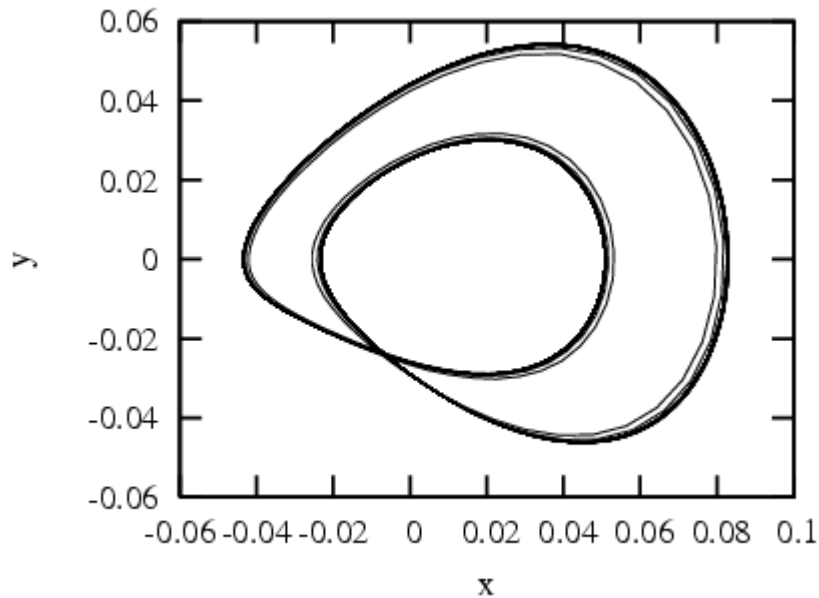
Nonlinear

$$\begin{aligned}\frac{dx}{dt} &= y \\ \frac{dy}{dt} &= \alpha x + \mu y + z \\ \frac{dz}{dt} &= -\beta y - pz - qy + syz\end{aligned}$$

1D model pulsations

Table 1. Parameters of the stellar pulsation model

	α	β	μ	p	q	s	σ	\bar{x}
periodic	-0.5	0.5	0.5	3.2	0.5	1.0	0.2	0
chaotic	-0.5	0.5	0.5	4.0	0.5	1.0	0.3	0



Quasiperiodic behaviour

Quasiperiodic signal:

$$x_{\text{quasi}}(t) = a_1 \sin(f_1 t) + a_2 \sin(f_2 t) + a_3 \sin(f_3 t)$$

Gaussian noise:

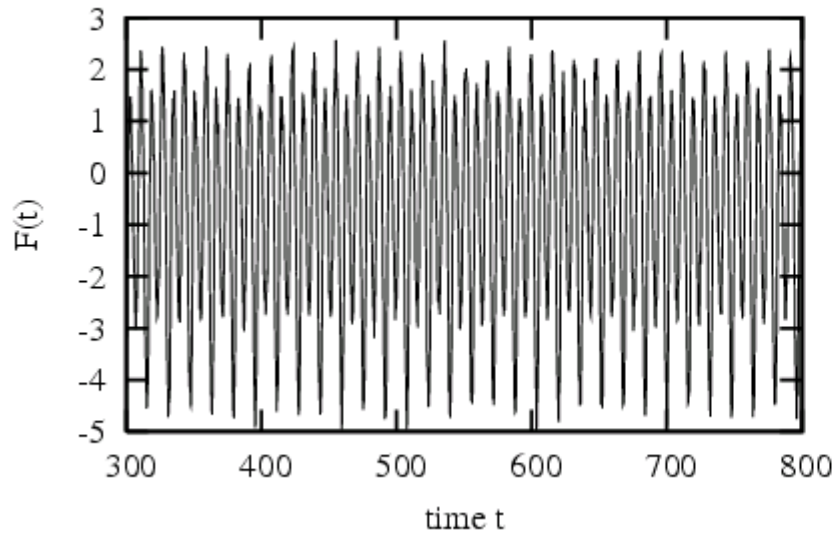
$$G(x) = \frac{1}{\sigma} \exp\left(-\frac{(x - \bar{x})^2}{2\sigma^2}\right)$$

Table 2. Parameters for the quasiperiodic signal

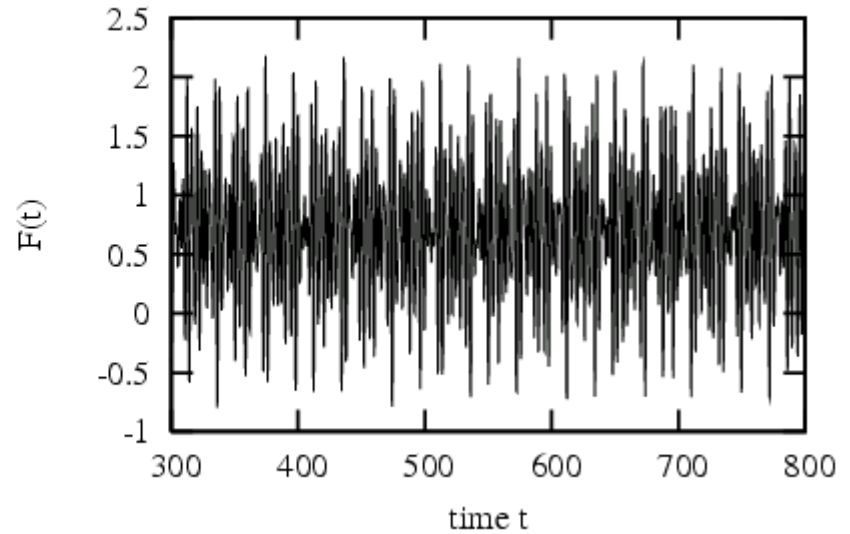
a_1	a_2	a_3	f_1	f_2	f_3	σ	\bar{x}
0.4	0.6	0.5	$\sqrt{5}$	$\sqrt{3}$	$\sqrt{2}$	0.1	0

Solution of equations

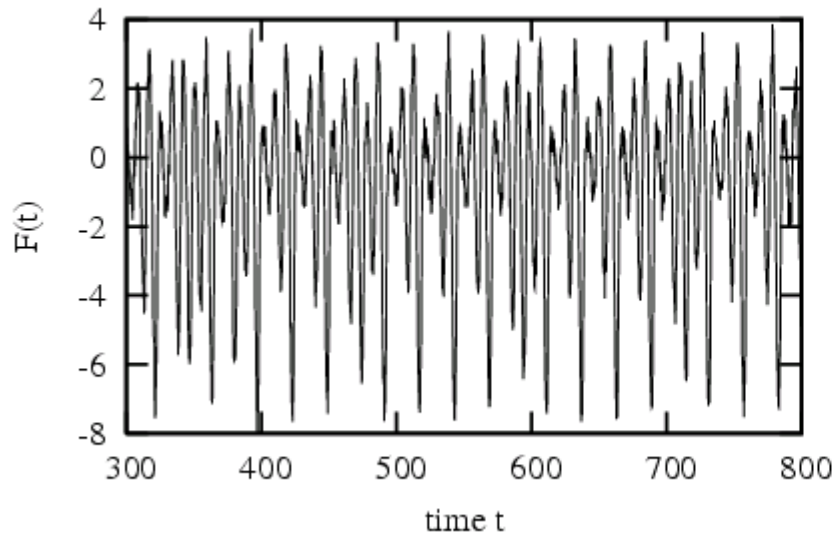
Periodic oscillation



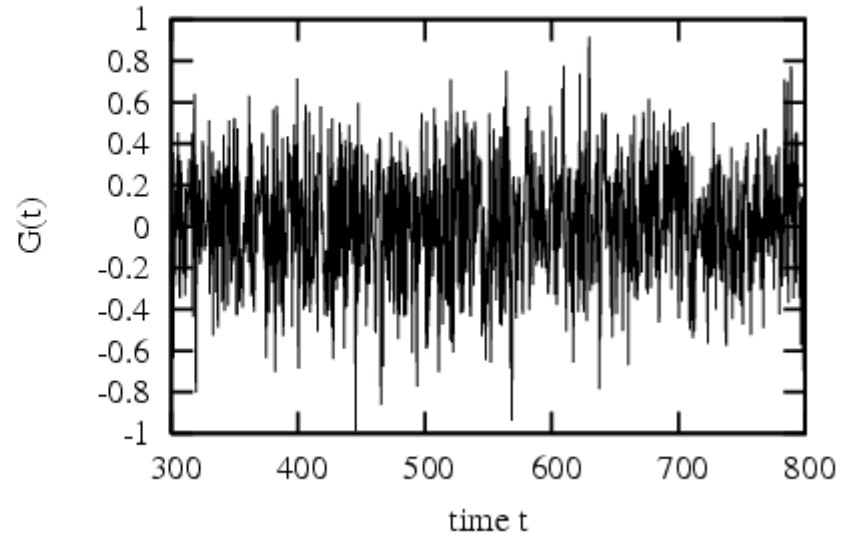
Quasiperiodic oscillation



Chaotic oscillation

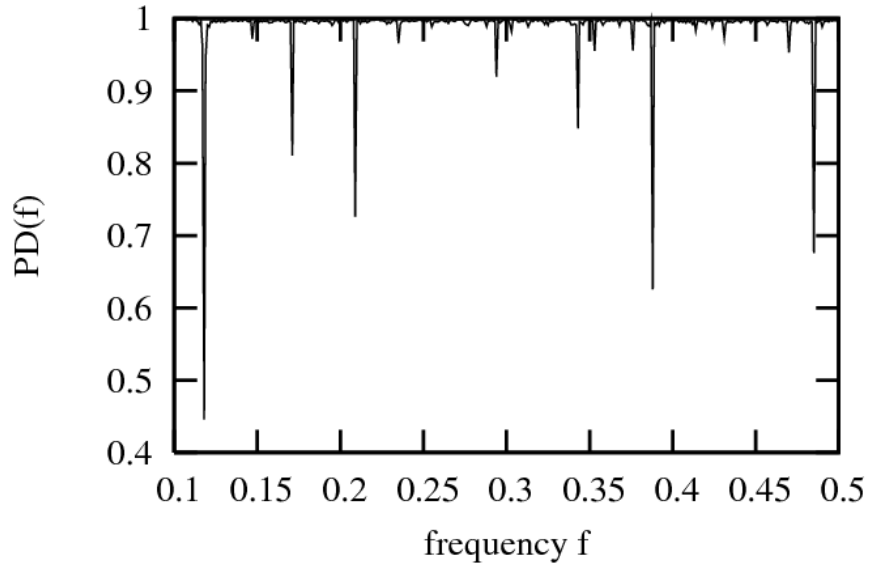


Gaussian noise

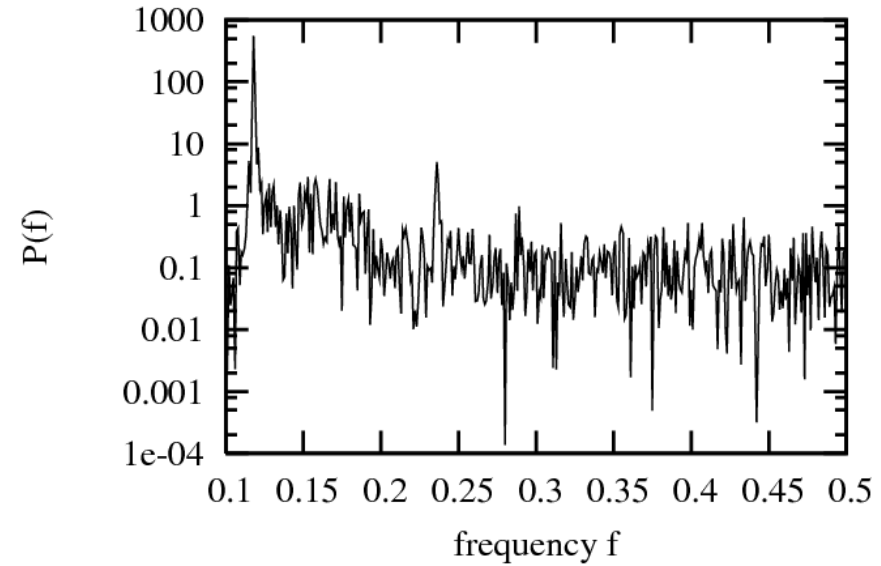


Period searches from time series

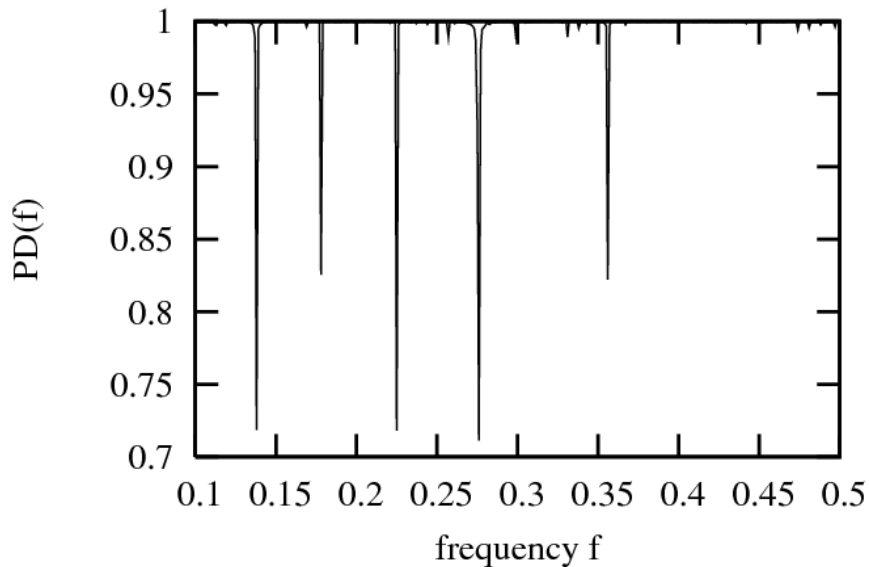
PDM period analysis



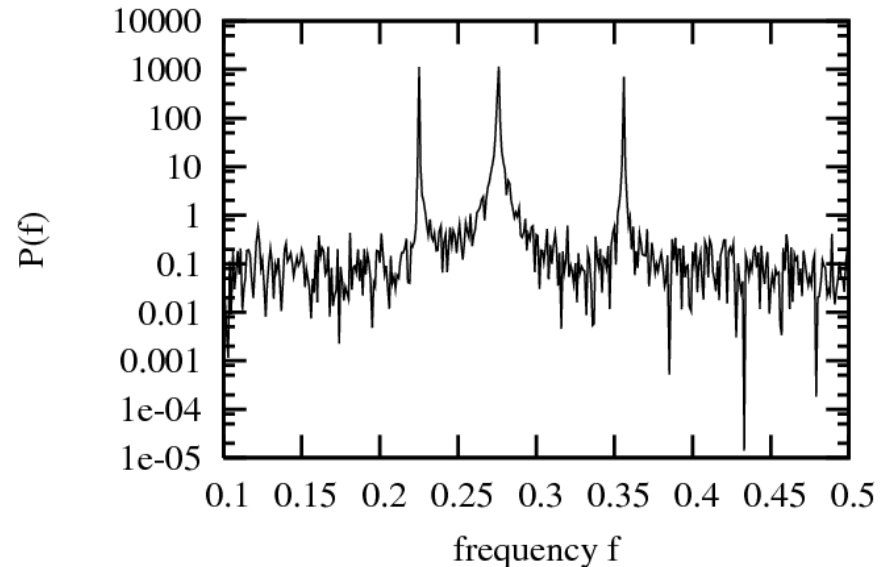
Power Spectrum



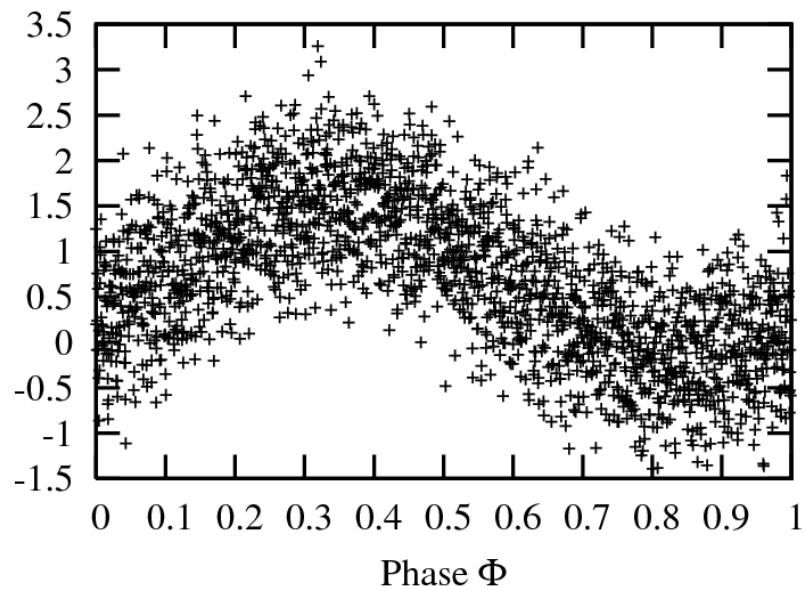
PDM period analysis



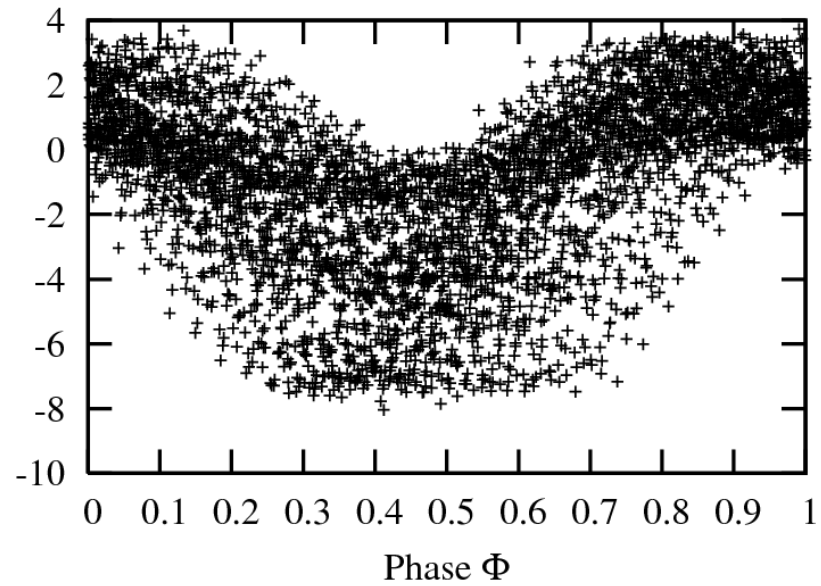
Power spectrum



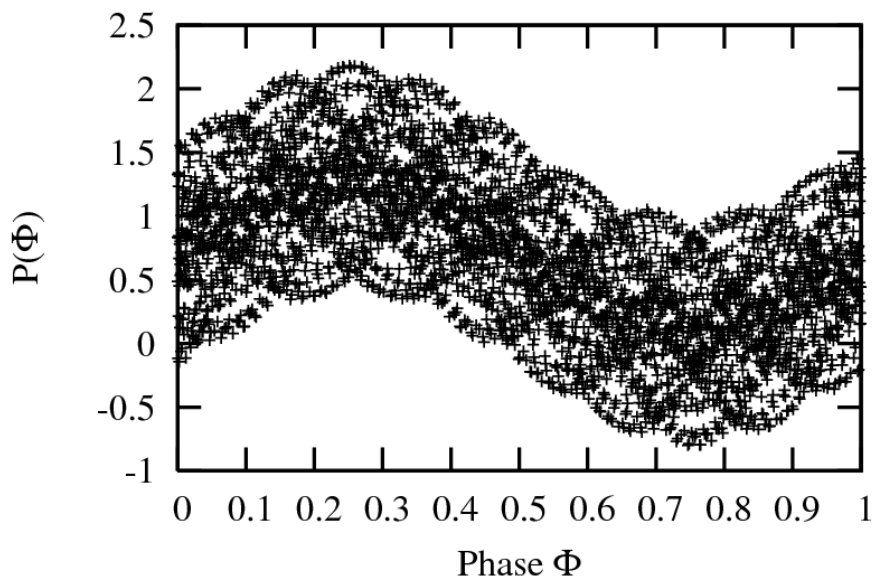
Phase diagram



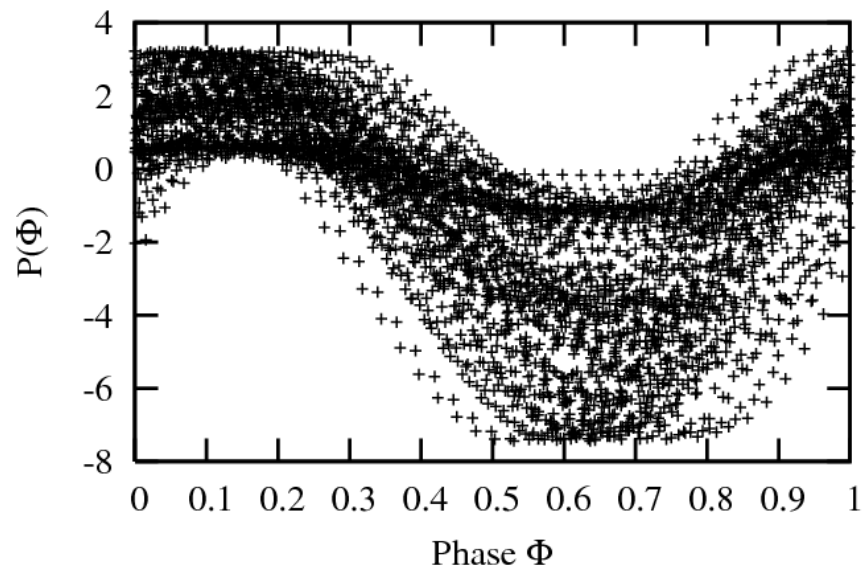
Phase diagram



Phase diagram



Phase diagram



Reconstruction of phase portrait

- Time series (equidistant)

$$X(t) = (x_1, x_2, x_3, \dots, x_N)$$

- R. Takens methods of time-delay τ

$$X_n = (x_1, x_{1+\tau}, x_{1+2\tau}, x_{1+3\tau}, \dots, x_{n+(m-1)\tau})$$

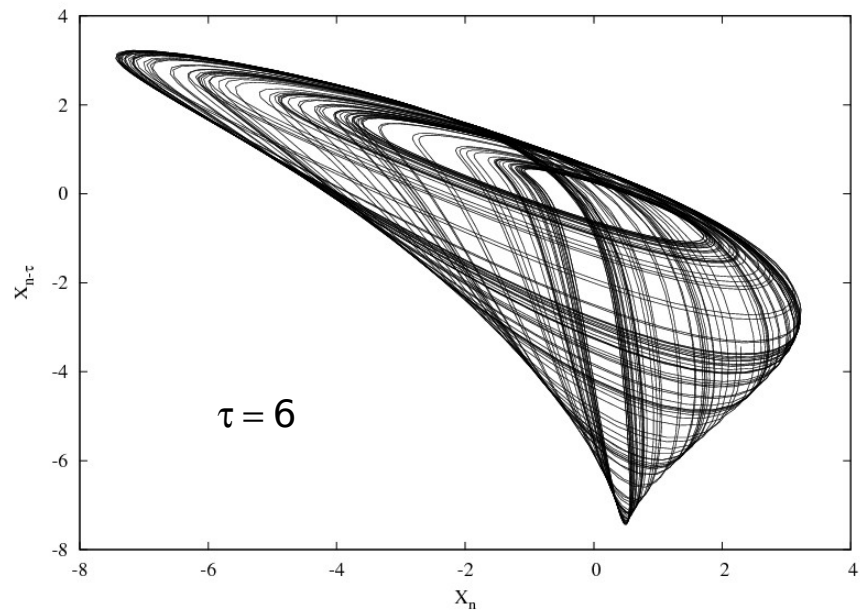
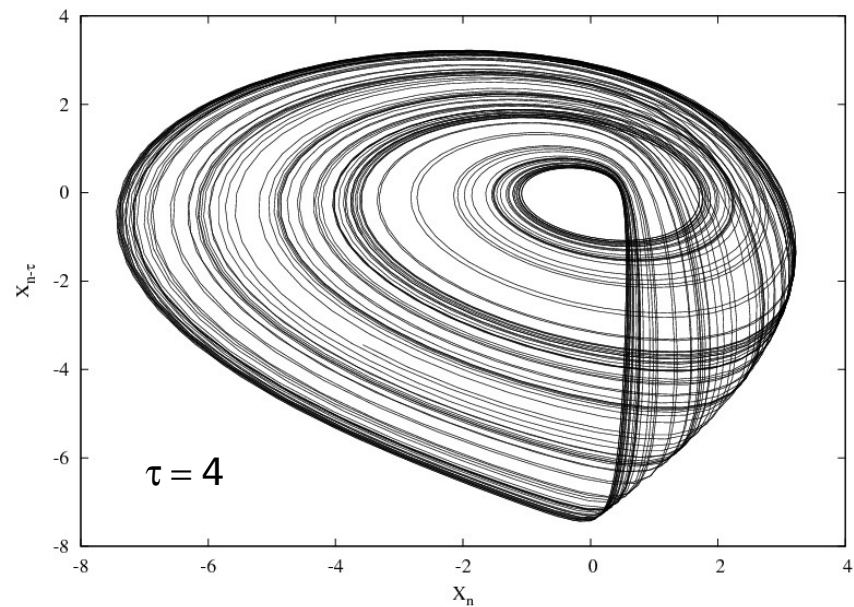
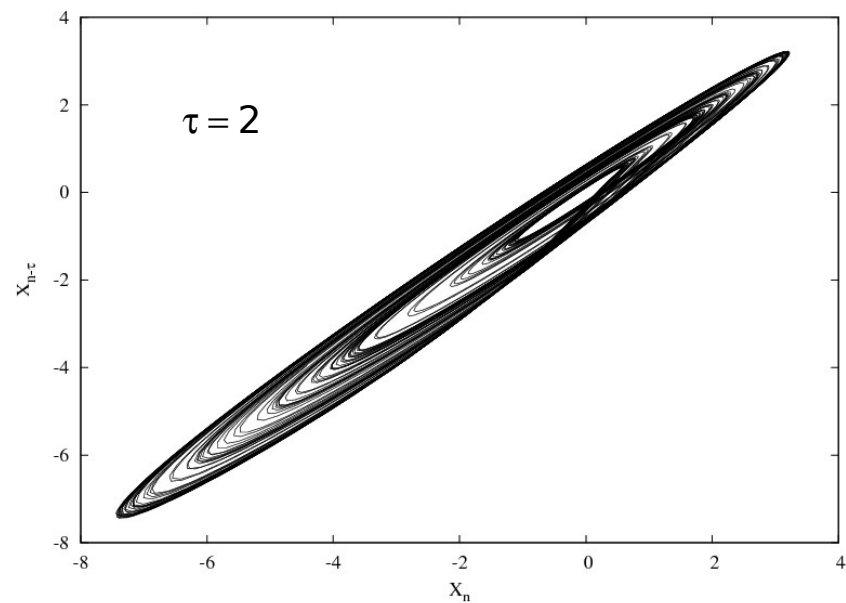
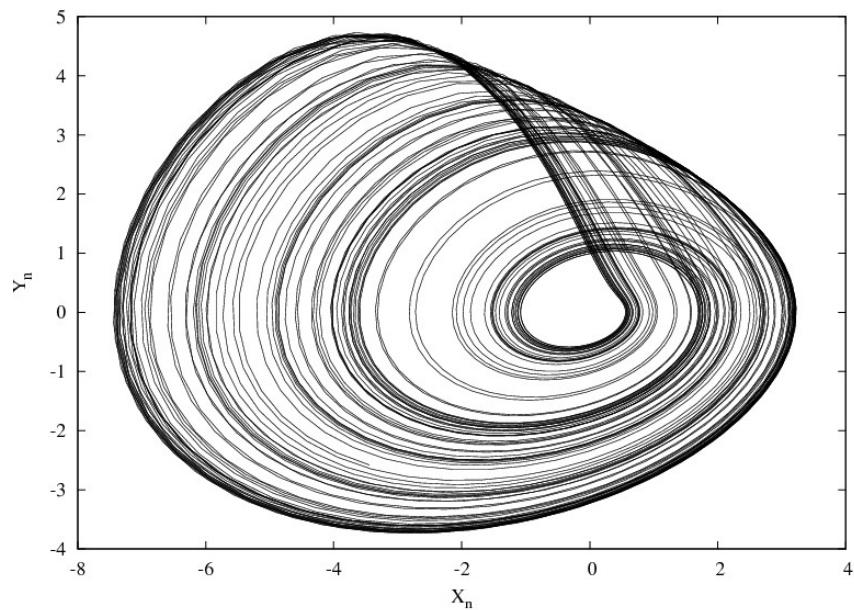
- Parameter of the method:

Time delay τ

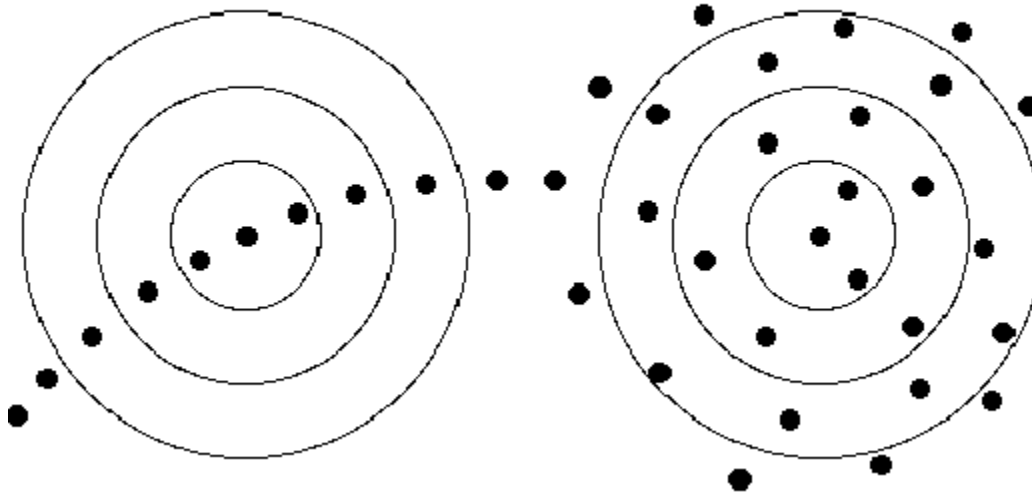
Dimension of reconstructed phase space m

- Reconstructed phase portrait have same topological invariants

Reconstruction of phase portrait

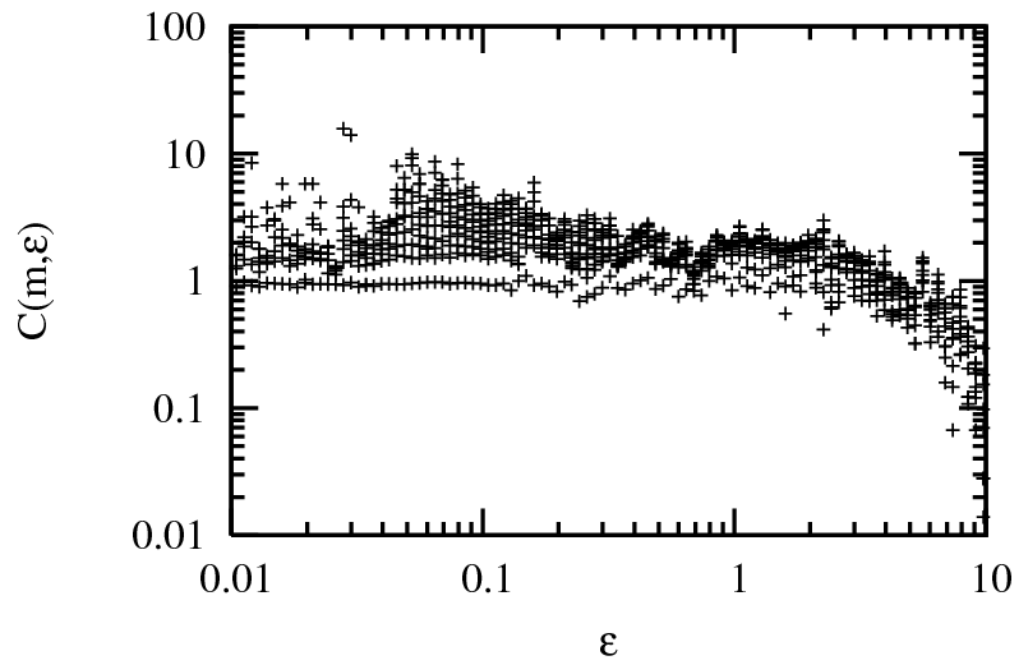


Correlation dimension



$$C(m, \epsilon) = \frac{2}{(N - n_{\min})(N - n_{\min} - 1)} \sum_{i=1}^N \sum_{j=i+1+n_{\min}}^N \theta(\epsilon - |s_i - s_j|)$$

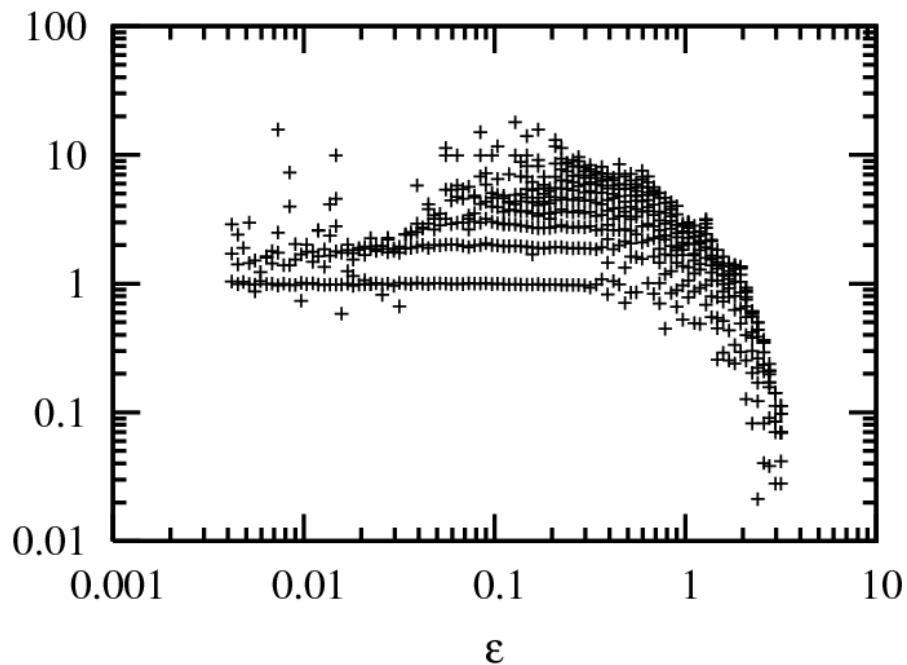
Correlation integral



Chaotic signal

$$C_{\text{kor}} = 1.9 \pm 0.8$$

Correlation integral



Quasiperiodic signal

Summary

- Simple nonlinear model of stellar pulsations lead to the chaotic oscillations
- It is difficult to distinguish between chaotic and quasiperiodic signal with using classical tools only
- When data contains important noise, quasiperiodic or chaotic character of the signal can be overlooked and simple periodic (false) solution with strongest frequency can be determined
- In order to definitely rule out presence of the chaos it is necessary to use the nonlinear time series analysis

Next step

- Chaotic behaviour does not mean unpredictable
- We can predict signal on certain time scale
- Reconstructed phase portrait can help us to derive simple physical model and find driving parameters
- With the knowledge of driving parameters and their values for different regime (chaotic or quasiperiodic) we can determine subgroups in HR diagram for such a star

First step: prediction of global behaviour

Mapping function:

$$F(\mathbf{X}) = \sum_k \mathbf{C}_k \mathbf{P}_k(\mathbf{X})$$

with polynomials:

$$P_{\mathbf{k}}(\mathbf{X}) = \sum_{\alpha} A_{\mathbf{k}}^{\alpha} (\mathbf{X}^{\alpha_1})^{k_1} (\mathbf{X}^{\alpha_2})^{k_2} \dots \quad \alpha = 1, 2, \dots, d$$

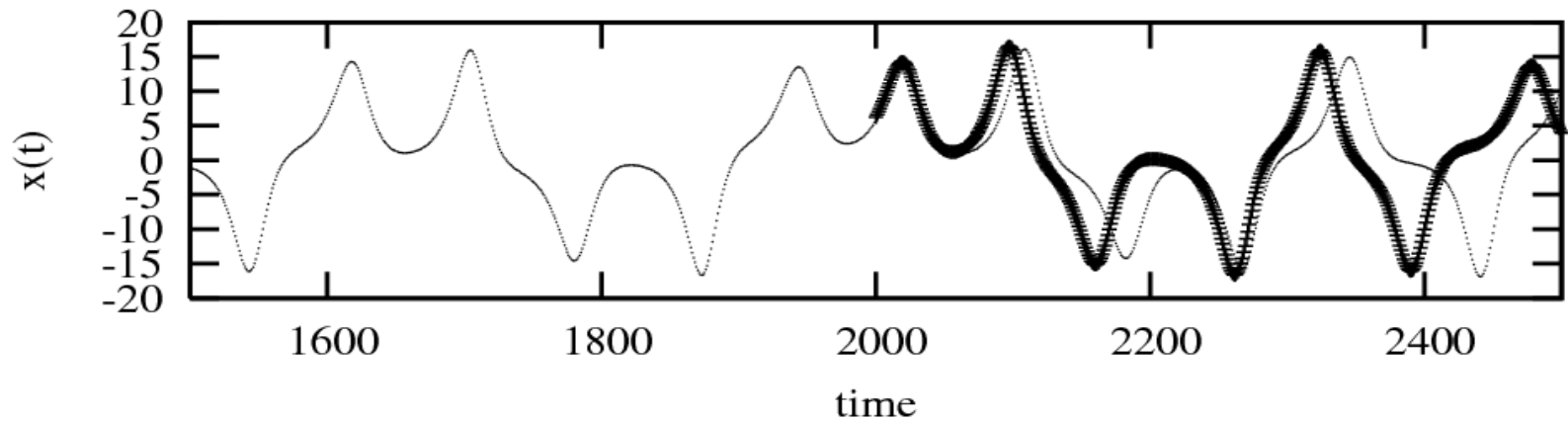
up to order

$$\sum_j k_j \leq p$$

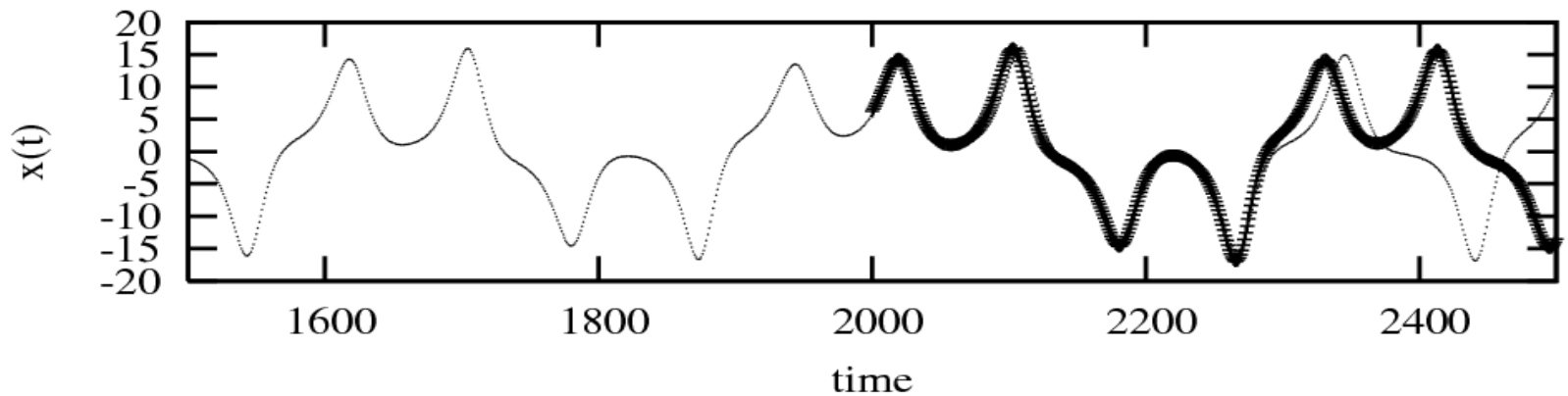
First step: prediction of global behaviour

$$B^K(\mathbf{X}) = \sum_{\alpha} V_{\alpha}^K(\mathbf{X}_{\alpha I})_{KI}(\mathbf{X}_{\alpha J})_{KJ} \dots$$

Global flow reconstruction



Global flow reconstruction



Reference:

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- Serre T., Kolláth Z., & Buchler, J.R., 1996 *Astronomy & Astrophysics*, 311, p. 833
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